Extremely red objects in the UKIDSS Ultra Deep Survey Early Data Release

Chris Simpson,1* Omar Almaini,2 Michele Cirasuolo,3 Jim Dunlop,3 Sébastien Foucaud,2 Paul Hirst,4 Rob Ivison,5,3 Mat Page,6 Steve Rawlings,7 Kaz Sekiguchi,8 Ian Smail9 and Mike Watson10

1Astrophysics Research Institute, Liverpool John Moores University, Twelve Quays House, Egerton Wharf, Birkenhead CH41 1LD
2School of Physics and Astronomy, University of Nottingham, University Park, Nottingham NG7 2RD
3Institute for Astronomy, University of Edinburgh, Royal Observatory, Blackford Hill, Edinburgh EH9 3HJ
4Joint Astronomy Centre, 660 N. A‘ohōkū Place, Hilo, HI 96720, USA
5Infrared Telescope Facility, University of Hawaii, 460 N. A‘ohōkū Place, Hilo, HI 96720, USA
6Department of Physics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford OX1 3RH
7Mullard Space Science Laboratory (MSSL), University College London, Holmbury St. Mary, Dorking, Surrey RH5 6NT
8Subaru Telescope, National Astronomical Observatory of Japan, 650 N. A‘ohōkū Place, Hilo, HI 96720, USA
9Institute for Computational Cosmology, Department of Physics, Durham University, South Road, Durham DH1 3LE
10Department of Physics and Astronomy, University of Leicester, Leicester LE1 7RH

Accepted 2006 August 18. Received 2006 July 31; in original form 2006 July 10

ABSTRACT
We construct a sample of extremely red objects (EROs) within the UKIDSS Ultra Deep Survey by combining the Early Data Release with optical data from the Subaru/XMM–Newton Deep Field. We find a total of 3715 objects over 2013 arcmin² with $R - K > 5.3$ and $K \leq 20.3$, which is a higher surface density than found by previous studies. This is partly due to our ability to use a small aperture in which to measure colours, but is also the result of a genuine overdensity of objects compared to other fields. We separate our sample into passively evolving and dusty star-forming galaxies using their $RJK$ colours and investigate their radio properties using a deep radio map. The dusty population has a higher fraction of individually detected radio sources and a higher mean radio flux density among the undetected objects, but the passive population has a higher fraction of bright radio sources, suggesting that active galactic nuclei are more prevalent among the passive ERO population.

Key words: surveys - galaxies: evolution - galaxies: starburst - infrared: galaxies.

1 INTRODUCTION
Extremely red objects (EROs) are variously defined on the basis of an optical–near-infrared colour, usually $R - K$ or $I - K$. These red colours can be due either to an evolved, passive stellar population at $z \gtrsim 1$, or the presence of large quantities of dust associated with starburst activity (although edge-on discs may also contribute significantly; Yan & Thompson 2003). While the former population reveals the stellar mass already assembled in galaxies as a result of star formation at early cosmic times, the latter population reveals the star formation at $z \sim 1$ which, due to its dust-enshrouded nature, may be missed by other surveys. As these two populations represent very different routes to the creation of present-day galaxies, there is a clear desire to reliably separate them and use them as tests of galaxy formation models.

Existing samples of EROs (for which we adopt the most common definition of $R - K > 5.3$) have been relatively small (typically no more than about 100 objects), due to the difficulty in obtaining near-infrared imaging data with the necessary depth and areal coverage (e.g. Daddi et al. 2000; Smail et al. 2002; Moustakas et al. 2004; Georgakakis et al. 2006). The arrival of infrared cameras with large fields of view, such as the Wide Field Infrared Camera (WFCAM; Casali et al., in preparation), has largely eliminated this problem, allowing samples of EROs to be constructed which are large enough to permit further separation by magnitude, colour, and other properties. Furthermore, by locating infrared surveys in well-observed patches of sky, it becomes possible to study the multiwavelength properties of EROs and learn more about their nature.

In this Letter, we combine near-infrared data from the UKIDSS Ultra Deep Survey Early Data Release (UDS EDR) (Dye 2006; Lawrence 2006) with optical data from the Subaru/XMM–Newton Deep Field (SXDF; Furusawa et al., in preparation) survey to construct a sample of EROs to faint magnitudes ($K > 20$) over an area
larger than half a square degree. The Ultra Deep Survey will eventu-
ally cover 0.8 deg$^2$ to 5$\sigma$ point source limits of $J = 25.0$, $H = 
24.0$ and $K = 23.0$ and has been sited in the SXDF because of the 
excellent deep multiwavelength data in this region. We determine 
photometric redshifts for all objects and separate them into dusty 
star-forming and passively evolving galaxies on the basis of their $RJK$
colours. We use deep radio data to determine the typical radio 
flux densities and star formation rates as a function of colour and 
magnitude. A more detailed analysis, including clustering and 
X-ray properties, will be undertaken after the UKIDSS DR1 release, 
which is expected to be a magnitude deeper than the present data, 
and will provide more reliable photometry. All magnitudes in this 
paper are on the Vega system. The SXDF catalogue is calibrated on 
the AB system, and we use the conversion $R_{\text{Vega}} = R_{\text{AB}} - 0.219$
(H. Furusawa, private communication). We adopt a cosmology with 
$H_0 = 70$ km s$^{-1}$ Mpc$^{-1}$ and $\Omega_{\Lambda} = 1 - \Omega_{\Lambda} = 0.3$.

2 THE ERO SAMPLE

We start with the improved version of the UDS EDR $K$-band cata-
logue produced by Foucaud et al. (2006), and determine the com-
pleteness for galaxies as a function of $K$-band magnitude. Galaxies 
with $K \approx 20$ have a median full width at half maximum FWHM = 
1.2 arcsec, and so we added objects of this size to the $K$-band WF-
CAM images and measured the fraction that were recovered by the 
extraction process. The 50 per cent completeness limit varies slightly 
between detectors, but averages to $K = 20.3$. We therefore elimi-
nated all objects from the catalogue of Foucaud et al. (2006) with 
(total) magnitudes fainter than this limit before proceeding further.
Our magnitude limit is brighter than that of Foucaud et al. because 
they calculate the completeness for unresolved sources. The optical 
data reach $R \approx < 26.7 (5\sigma, \phi 2$ arcsec).

2.1 Optical-infrared matching

In order to facilitate matching of the near-infrared catalogue with 
objects in the optical imaging, objects in various regions of sky 
were next removed from the $K \leq 20.3$ catalogue. Only the central 
region of the UDS image, with uniform noise, was included in the 
matching process. Objects were eliminated whose UDS coordinates 
placed them outside the optical imaging region, or within 2 arcsec 
of the edge, to ensure that the photometric aperture of any optical 
match was entirely within the image. Finally, objects which lay 
within the halo or CCD bleed of a bright star were removed from 
the catalogue. The remaining effective area is 2013 arcmin$^2$.

The UDS catalogue was first matched to the catalogues of Furus-
awa et al. (in preparation), with the closest match within 1 arcsec 
being used. However, visual inspection of the unmatched objects re-
vealed that many pairs of objects remain blended in the optical cata-
logues, so the UDS catalogue was instead matched to the catalogues 
produced by Simpson et al. (2006) with a more severe deblending 
threshold. The $i'$-band catalogue was used first (as the $i$-band data 
are the deepest), followed by the $z'$, $R$, and $B$-band catalogues. This 
reduced the number of unmatched objects by about half, to 426.
These were all inspected by eye, and the vast majority (301) were 
discovered to be artifacts such as cross-talk, satellite trails, or array 
defects not removed in the reduction process (see Dye 2006). These 
were deleted from the catalogue and not considered further. Most of 
the remaining 125 sources are still blended in the optical catalogue, 
although some lack any optical counterpart. Optical photometry was 
determined for these objects directly from the Suprime-Cam images 
using the UDS catalogue coordinates. A total of 443 objects have 
measured $R$-band aperture magnitudes fainter than the $5\sigma$ limit of 
the optical data.

In what follows, colours are measured in a 2-arcsec apen-
ture, while magnitudes are total magnitudes from SExtractor’s 
$\text{MAG\_AUTO}$ parameter (Bertin & Arnouts 1996). The $K$-band image 
has a narrower point spread function (psf) than the other images 
(FWHM = 0.69 arcsec, compared to 0.81 arcsec), and we address 
this point below.

2.2 Star–galaxy separation

We classify objects as stars or galaxies by the same method as Simp-
son & Rawlings (2002), using a combination of $J - K$ colour and 
Kron (1980) radius, $r_{\text{\_2}}$. There is a clear bimodality in this plot 
(Fig. 1) because galaxies rapidly become redder than stars with in-
creasing redshift, while those with similar colours to stars are nearby 
and hence clearly extended. While there was a clear separation be-
tween the two groups in the small sample of Simpson & Rawlings 
(2002), the much larger number of objects in the UDS EDR means 
that many are scattered into ‘no man’s land’, and we determine the 
locus of the minimum surface density of objects in Fig. 1 as our 
separation criterion. All objects with $r_{\text{\_2}} < 0.41 - 0.1(J - K)$ are 
classified as stars. However, the combination of depth and area in the 
UDS will result in the detection of a number of very late-type stars 
with extremely red $J - K$ colours, and these will not be excluded by 
this single criterion. We therefore also classify as stars any objects 
with $r_{\text{\_2}} < 0.3$ and $K < 19.5$ (morphological classification is not 
possible fainter than this).

Our galaxy sample will suffer contamination from L dwarf stars, 
which have $J - K > 1$ (e.g. Hewett et al. 2006). Objects with $J - K > 1.3$ and $K > 19.5$ will be classified as galaxies, and 
these red colours are possessed by stars of spectral types L1–L8 
(table 9 of Knapp et al. 2004), whose surface density to 
$K < 14.5$ is $0.038$ deg$^{-2}$ (Reid et al. 1999). If these are isotropically distributed 
on the sky (so the number expected is proportional to the Euclidian 
volume sampled), we expect $\sim 60$ objects with $K \leq 20.3$ in our 
survey area. These are therefore a small contaminant.

2.3 Number counts

To construct our ERO sample, we first investigate the effect of the 
superior $K$-band image quality by Gaussian-smoothing objects with

Figure 1. Plot of $r_{\text{\_2}}$ (Kron 1980) against $J - K$ colour for all objects with $K < 20.3$. The solid line shows the criterion used to separate stars from galaxies.

R − K > 5.3 to the same FWHM as the other images and measuring the flux loss in a 2-arcsec aperture. We find a median reduction in brightness of 0.04 mag, and apply this reduction to all K-band magnitudes in the UDS EDR catalogue before constructing our ERO sample. In total we find 3715 (5128; 1464) objects with R − K > 5.3 (R − K > 5; R − K > 6).

In Fig. 2 we compare our galaxy and ERO number counts with determinations from the literature. Our total galaxy number counts agree very well, except at the brightest magnitudes where there is a slight deficit due to the fact that we have excluded regions with very extended galaxies from our analysis. Our ERO counts are, however, ~60 per cent higher than those of Daddi et al. (2000). This could be due in part to the different R-band filters used in these studies, as Daddi et al.'s filter has a long red tail with its 10 per cent point at 8200 Å, while the Suprime-Cam filter has a sharper long-wavelength cut-off with a 10 per cent transmission point at 7200 Å. Consequently, objects with a pronounced (4000-Å) break between these two wavelengths will appear fainter in the Suprime-Cam images, and hence redder in R − K. However, we note that we also find a much larger surface density of R − K > 5.3 objects than did Smail et al. (2002), who used the same instrument/filter combination for their optical imaging. This may be partly due to the small aperture we use to determine colours, since contamination from nearby objects (which is obviously more likely for larger apertures) will almost certainly make the aperture colour bluer. We find that the number of EROS is decreased by 15−20 per cent (with a larger reduction for redder objects) if we use the colours measured in a 4-arcsec aperture. Finally, we note that our surface density is similar to that of Miyazaki et al. (2003) who studied a small region within the SXDF, suggesting that this region is overdense at z ∼ 1. For objects as clustered as EROS (r0 ≈ 12 Mpc; Daddi et al. 2001), the effect of cosmic variance can be large, and we estimate σ1 ∼ 0.2 using the prescription of Somerville et al. (2004). We stress that the depth of the R-band data (the 3σ limit in a 2-arcsec aperture is R ≈ 27.4; Furusawa et al., in preparation) means that photometric errors will not scatter a significant number of the more numerous bluer objects into the ERO class.

3 NATURE OF THE ERO POPULATION

Although the ERO population is composed of both passively evolving red galaxies and dusty starburst galaxies, there does not appear to be a bimodality in any photometric property which allows these two classes to be readily separated. Pozzetti & Mannucci (2000) used J-band photometry to attempt to separate the classes as heavily-reddened objects will not show the pronounced 4000-Å break apparent in the spectra of old galaxies, and hence will have bluer R − J and redder J − K colours. Smail et al. (2002) used a more sophisticated method of spectral template fitting with the code HYPERZ (Bolzonella, Miralles & Pelló 2000) and classified objects according to whether the best-fitting model had a low or high extinction, finding that this produced broadly the same classification as the simpler method. Both methods rely on the accuracy of the photometry around λrest ∼ 4000 Å, and our ability to classify EROS is therefore limited by the quality of the J-band photometry. Because the dividing line between the two classes of object is at J − K ≲ 2 and the J-band photometry limit is J = 21.6 (5σ, p2 arcsec), a classification is likely to be reliable only for the 1922 EROS with K ≤ 19.5 and we do not believe a sophisticated analysis is warranted by the present near-infrared data quality. We therefore opt to use Pozzetti & Mannucci’s (2000) simple RIK colour criterion to separate our sample into passive and dusty galaxies. We find that 62 per cent of K ≤ 19.5 EROS are classified as passively evolving, although this fraction increases when a redder colour cut is used, rising to 73 per cent (77 per cent) of objects with R − K > 6 (R − K > 6.5). These fractions are broadly consistent with, although perhaps slightly larger than, those determined from previous studies.

We have computed photometric redshifts for all members of our ERO sample, in order that we can derive physical quantities from observed fluxes, by fitting synthetic galaxy templates to the BVRiJK photometry, plus the two short-wavelength IRAC bands (from the SWIRE survey of Lonsdale et al. 2003) where available, using a modified version of the HYPERZ code (see Cirasuolo et al. 2006, for a detailed description). We show the results of this analysis by plotting the distribution for various subsamples in Fig. 3. Objects whose best template fits have χ2 > 20.09 (corresponding to an unacceptable fit at 99 per cent confidence) are not plotted.

Fig. 3 also reveals the predictable results that the reddest objects are typically at higher redshift, as are the faintest objects. The redshift distribution for passive galaxies also appears to be slightly narrower and more strongly peaked than that for dusty ones. This is due to the passage of the 4000-Å break through the filters, as galaxies will only be red in R − K if the break is beyond the blue end of the K filter, while the Pozzetti & Mannucci (2000) criterion breaks down at z > 2 and classifies all EROS as dusty. We caution against a more detailed analysis of the different N(z) distributions given the significant difference between the photometric accuracy of the optical and near-infrared data, even for EROS. The deeper JK data from UKIDSS DR1 will provide better photometry, and hence a more reliable photometric analysis, for the present sample, as well as allowing a basic study of the fainter (K ∼ 21) ERO population.

4 RADIO PROPERTIES

Simpson et al. (2006) have presented a 1.4-GHz radio image of the SXDF taken with the NRAO Very Large Array which reaches an rms noise level of 12 μJy beam−1 in the central regions. At the median redshift of our sample, z ≈ 1.3, a star formation rate (SFR) for M > 5 M⊙ stars of 10 M⊙ yr−1 produces a flux density of 5 μJy (Condon 1992). Although this is obviously below the detection threshold for any individual source, we can make statistical detections of relatively modest SFRs in samples of ≥100 objects as the 5 × 4 arcsec2 beam will encompass all radio emission from the host galaxy. We
note that starburst galaxies will be unresolved at the resolution of this radio image.

We study the mean radio properties of our sample in two ways. First, we produce small postage stamps from the radio map of Simpson et al. (2006) at our objects’ positions and sum these after weighting each by the inverse of the pixel noise variance at its position. We exclude all objects where any of the nine pixels (the scale is 1.25 arcsec pixel$^{-1}$) in the radio image closest to the source position has a signal-to-noise ratio in excess of 3, so that the mean is not biased by a few objects associated with bright radio sources. We then fit an elliptical Gaussian profile to this stacked image using Ivan Busko’s N2GAUSSFIT task in IRAF, and calculating the flux density from the fit using the prescription of Condon (1997). For the second method, we construct a histogram of pixel values from the radio map at the positions of objects in our sample, and fit a Gaussian to the pixels after they have been iteratively sigma-clipped to exclude values more than two standard deviations from the mean, thus automatically removing bright outliers. We refer to these two methods as the stacking and histogram methods, respectively. Because the histogram method fits a Gaussian to the convolution of true source fluxes with the near-Gaussian pixel noise distribution, it is most reliable if the source flux density distribution is either itself Gaussian, or has a dispersion much less than that of the overall pixel noise. We prefer to use the image stacking method to derive mean radio flux densities, but confirm that the results from the histogram method are consistent in all cases.

To determine the likelihood of spurious detections in our image stacks, we ran Monte Carlo tests by producing 1000 lists of between 100 and 1000 sources each located at random positions within the unmasked region of the SXDF/UDS images. The mean flux density at the positions in each list was then calculated by both the stacking and histogram methods. Because the stacked radio images are essentially signal-free, the Gaussian fitting technique was highly unstable and instead we measured the signal in a 5 × 5 pixel box in the centre of the image and corrected this for the 20 per cent flux loss expected for the beam size. Both methods produced a mean flux density of $-0.2 \mu$Jy, and we have therefore corrected our measured flux densities for this offset. The standard deviations on the mean for a sample of 100 sources were 0.9 and 2.1 $\mu$Jy for the stacking and histogram methods, respectively, and reduced with sample size $N$ approximately as $N^{-1/2}$.

Table 1 shows the results of the stacking analysis for 18 ERO subsamples (separated by magnitude, colour and classification), while Fig. 4 shows the stacking and histogram analyses for one subsample, indicating the reality of the statistical detection. From our analysis, galaxies classified as ‘dusty’ typically have twice the mean radio flux density of ‘passive’ galaxies, and the fraction of individually detected objects is about twice as high (except in the faintest magnitude bin, where we have argued that the classification is unreliable). Smail et al. (2002) also find that most EROs with flux densities of a few tens of $\mu$Jy are dusty. However, while a smaller fraction of

| Colour threshold | Magnitude range | $\tau_{\text{phot}}$ | $N_{\text{stack}}$ | ‘Passive’ EROs | $\overline{S_{1.4}}$ (μJy) | SFR ($M_\odot$ yr$^{-1}$) | $\tau_{\text{phot}}$ | $N_{\text{stack}}$ | ‘Dusty’ EROs | $\overline{S_{1.4}}$ (μJy) | SFR ($M_\odot$ yr$^{-1}$) |
|------------------|-----------------|---------------------|------------------|-------------------|--------------------------|----------------|-----------------|------------------------|--------------------------|----------------|----------------|---------------------------|
| $R - K > 5.3$    | $K \leq 18.5$   | 1.05                | 220              | 54                | 6.7 ± 1.3               | 8.5 ± 1.7 | 0.96             | 109                    | 55                       | 15.0 ± 1.0 | 15.3 ± 1.0 |                           |
|                 | $18.5 < K \leq 19.5$ | 1.20               | 819              | 104               | 4.2 ± 0.6               | 7.2 ± 1.1 | 1.25             | 463                    | 98                       | 10.3 ± 0.5 | 19.7 ± 0.9 |                           |
|                 | $19.5 < K \leq 20.3$ | 1.37               | 832              | 78                | 3.4 ± 0.4               | 8.1 ± 0.9 | 1.46             | 800                    | 83                       | 5.3 ± 0.7   | 14.5 ± 2.0 |                           |
| $R - K > 6.0$    | $K \leq 18.5$   | 1.24                | 64               | 12                | 9.4 ± 1.7               | 17.6 ± 3.1 | 1.18            | 14                     | 8                        | 11.7 ± 1.9 | 19.5 ± 3.2 |                           |
|                 | $18.5 < K \leq 19.5$ | 1.36               | 351              | 55                | 5.3 ± 0.4               | 12.3 ± 0.9 | 1.32             | 131                    | 29                       | 9.4 ± 0.7   | 19.9 ± 1.5 |                           |
|                 | $19.5 < K \leq 20.3$ | 1.44               | 414              | 34                | 3.1 ± 0.6               | 8.2 ± 1.5  | 1.60             | 321                    | 31                       | 4.4 ± 1.2   | 15.0 ± 4.0 |                           |
| $R - K > 6.5$    | $K \leq 18.5$   | 1.35                | 19               | 2                 | 12.1 ± 2.0              | 27.7 ± 4.6 | 1.53            | 6                      | 1                        | 10.5 ± 4.8* | 32.3 ± 14.7 |                           |
|                 | $18.5 < K \leq 19.5$ | 1.47               | 131              | 25                | 5.8 ± 0.7               | 16.3 ± 1.9 | 1.68             | 40                     | 7                        | 10.2 ± 2.0 | 39.2 ± 7.6 |                           |
|                 | $19.5 < K \leq 20.3$ | 1.64               | 177              | 16                | 3.4 ± 0.8               | 12.2 ± 2.7 | 1.81             | 124                    | 10                       | 4.8 ± 1.2   | 21.6 ± 5.5 |                           |

*aEllipse parameters fixed during fit due to high noise.*
passive galaxies are detected individually, those that are tend to be brighter: in the $18.5 < K \leq 19.5$ bin, 28 (13) of 104 individually detected passive EROs have $S_{1.4} > 100 \mu$Jy ($S_{1.4} > 300 \mu$Jy), while only 10 (four) of 98 dusty EROs are as bright (for reference, the Fanaroff & Riley (1974) break is at $S_{1.4} \approx 6 \mu$Jy at $z = 1.3$). This suggests that a few per cent of the passive ERO population may be active galaxies (cf. Alexander et al. 2002), although such objects can also be found in the dusty population (e.g. EDXS N2.21; Willott et al. 2003). Among the dusty population, we find that the SFR is uncorrelated with magnitude, but may be correlated with colour.

Our mean radio flux densities are consistent with those determined by Georgakakis et al. (2006), although our larger sample size produces much smaller uncertainties. We find that dusty galaxies typically have about twice the 1.4-GHz flux density of passive galaxies with the same colours and magnitudes, and hence have about twice the star formation rate because the median redshifts of identically selected passive and dusty EROs are similar (in computing the median redshifts for our subsamples, we ignore cases where the best-fitting photometric redshift has $\chi^2 > 20.09$). However, the true difference in star formation rates is likely to be larger than this because the individually detected dusty objects are likely to be the high end of a broad distribution of star formation rates whereas the individually detected passive objects are likely to be active galactic nuclei. As a final point, we note that the sample of 5834 non-ERO galaxies with $18.5 < K \leq 19.5$ and $R - K \leq 5.3$ has 440 objects detected at $>3\sigma$ significance, while the remaining sources have a mean flux density of $S_{1.4} = 5.3 \pm 0.2 \mu$Jy, consistent with the value derived for the passive ERO population in this magnitude range.

5 SUMMARY

We have described the construction of a sample of extremely red objects (EROs) within the Subaru/XMM-Newton Deep Field (SXDF) and UKIDSS Ultra Deep Survey (UDS) which is much larger than previous samples. We have measured a higher surface density than found in other fields, but comparable to that measured in a small sub-region of our field, suggesting the SXDF samples a cosmic overdensity at $z \sim 1$. After using RJK colours to separate the sample into passively evolving (62 per cent of EROs with $K \leq 19.5$) and dusty star-forming galaxies, we have used a deep radio map to study the radio properties of these two populations. Both populations are detected at high significance in a stacked radio image, with the dusty sample having a mean flux density (and derived mean star formation rate) equal to approximately twice that of the passive sample. The mean radio flux density of the passive sample is consistent with that of non-ERO galaxies with the same magnitudes.

ACKNOWLEDGMENTS

This paper is partially based on data collected at the Subaru Telescope, which is operated by the National Astronomical Observatory of Japan. The United Kingdom Infrared Telescope is operated by the Joint Astronomy Centre on behalf of the UK Particle Physics and Astronomy Research Council (PPARC). The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. We thank the staff of these facilities for making these observations possible, and the anonymous referee for helpful suggestions. We gratefully acknowledge financial support from the PPARC (CS, MC, SF, SR) and the Royal Society (OA, IS).

REFERENCES

Miyazaki M. et al., 2003, PASJ, 55, 1079
Moustakas L. A. et al., 2004, Apl, 600, L131
Reid I. N. et al., 1999, Apl, 521, 613
Yan L., Thompson D., 2003, Apl, 586, 765

This paper has been typeset from a TeX document prepared by the author.