ABSTRACT

We present a new technique for identifying stellar mass black holes in low mass X-ray binaries (LMXBs), and apply it to XMM-Newton observations of M31. We examine X-ray time series variability seeking power density spectra (PDS) typical of LMXBs accreting at a low accretion rate (which we refer to as Type A PDS); these are very similar for black hole and neutron star LMXBs. Galactic neutron star LMXBs exhibit Type A PDS at low luminosities ($10^{36} - 10^{37}$ erg/s) while black hole LMXBs can exhibit them at luminosities $>10^{38}$ erg s$^{-1}$. We propose that Type A PDS are confined to luminosities below a critical fraction of the Eddington limit, $l_c$, that is constant for all LMXBs; we have examined a sample of black hole and neutron star LMXBs and find they are all consistent with $l_c = 0.10 \pm 0.04$ in the 0.3–10 keV band. We present luminosity and PDS data from 167 observations of X-ray binaries in M31 that provide strong support for our hypothesis. Since the theoretical maximum mass for a neutron star is 3.1 $M_\odot$, we therefore assert that any LMXB that exhibits a Type A PDS at a 0.3–10 keV luminosity greater than $4 \times 10^{37}$ erg s$^{-1}$ is likely to contain a black hole primary. We have found eleven new black hole candidates in M31 using this method. We focus on XMM-Newton observations of RX J0042.4+4112, an X-ray source in M31 and find the mass of the primary to be $7 \pm 2$ $M_\odot$, if our assumptions are correct. Furthermore, RX J0042.4+4112 is consistently bright in 40 observations made over 23 years, and is likely to be a persistently bright LMXB; by contrast all known Galactic black hole LMXBs are transient. Hence our method may be used to find black holes in known, persistently bright Galactic LMXBs and also in LMXBs in other galaxies.

Subject headings: X-rays: general — X-rays: binaries — Galaxies: individual: M31 — black hole physics— Methods: data analysis

1. Introduction

Van der Klis (1994) showed that the power density spectra (PDS) of low mass X-ray binaries (LMXBs) with neutron star primaries at low accretion rates are strikingly similar to those of black hole LMXBs in their low accretion rate states. The fractional rms variability is high ($\sim 30-50\%$) and the PDS are well described by a broken power law that changes in spectral index, $\gamma$, from $\sim 0$ to $\sim 1$ at frequencies higher than a certain break frequency; the break occurs at $0.01-1$ Hz. We will refer to these as Type A PDS. At higher accretion rates, the rms variability is only a few percent, and the PDS are characterised by a power law with $\gamma \sim 1-1.5$. We will refer to these as Type B PDS. Van der Klis (1994) proposed that the transition between Type A and Type B PDS occurs when the accretion rate exceeds a critical fraction of the Eddington limit ($f_c$) that is constant for all LMXBs. He suggested $f_c \sim 1\%$, as an order of magnitude estimate.

We propose here a new diagnostic for identifying stellar-mass black holes in LMXBs. Since we cannot observe $\dot{m}$ directly, we must use the luminosity to trace the evolution of the PDS with $l$. We define $l$ as $L/L_{Edd}$, where $L$ is the luminosity and $L_{Edd}$ is the Eddington limit, and $l_c = L_c/L_{Edd}$ so that Type A PDS are exhibited by LMXBs when $l < l_c$ and Type B PDS are exhibit-
Fig. 1.— Example PDS from XMM-Newton observations of the central region of M31; the axes are log-scaled and the PDS are Leahy normalised, so that Poisson noise has a power of 2. The left panel shows a Type A PDS: the best fit power law to the PDS has an unacceptable $\chi^2$ of 51/14, and the fractional r.m.s. variability of the lightcurve is 30%. The right panel shows a Type B PDS: its shape is well described by a power law ($\chi^2$/dof = 9/14), and the fractional r.m.s variability of the lightcurve is 5%.

2. Obtaining an empirical value for $l_c$ in the 0.3–10 keV band

In practise, we do not observe the full bolometric luminosity of an X-ray source and can only obtain its luminosity in a given energy band. We have been using the XMM-Newton and Chandra X-ray observatories to identify black hole LMXBs in external galaxies, and these have an energy range of $\sim$0.3–10 keV. We therefore estimate $l_c$ in the 0.3–10 keV band, i.e. $L_{c}^{0.3–10\text{keV}}/L_{\text{Edd}}$, since it is directly applicable to our observations. The energy spectra of neutron star and black hole LMXBs are similar at low accretion rate; hence $l_c$ should scale to different energy bands in a similar way for all LMXBs.

2.1. Estimating $l_c$ from XMM-Newton observations of 14 globular cluster LMXBs in M31

Of the 63 X-ray source that we studied in the central region of M31, before focusing on RXJ0042.4+4112, an X-ray source in the vicinity of M31 that appears to exhibit both Type A and Type B PDS, allowing us to estimate its mass.
globular clusters by Kong et al. (2002). There are thirteen bright X-ray sources in Galactic globular clusters, and twelve of these have been identified as neutron star LMXBs, while the thirteenth has not been classified (in’t Zand et al. 2004, and references within). Hence the 14 globular cluster X-ray sources in our sample are expected to be LMXBs containing \( \sim 1.4 \, M_\odot \) neutron stars. Bo 153 and Mita 299 are likely to be multiple bright X-ray sources, but are possible black holes.

Each X-ray source is consistent with the hypothesis that Type A PDS are exhibited at lower luminosities than Type B PDS, within errors. Furthermore, twelve are consistent with \( 1.0 \leq (L_c / 10^{37} \, \text{erg s}^{-1}) \leq 2.6 \), assuming a distance to M31 of 760 kpc (van den Bergh 2000); this corresponds to \( L_c = 0.10 \pm 0.04 \) if the primaries in these twelve sources are 1.4 \( M_\odot \) neutron stars. Bo 153 and Mita 299 exhibit Type A PDS at luminosities of \( (10.3 \pm 0.8) \) and \( (6.7 \pm 0.6) \times 10^{37} \, \text{erg s}^{-1} \) respectively. These systems would be consistent with our hypothesis if they contained black hole binaries, or if they were composed of two or more bright X-ray sources (like M15, see White & Angelini 2001) or multiple faint X-ray sources, in the globular clusters (see e.g. Heinke et al. 2003).

2.2. \( l_c \) for a Galactic neutron star LMXB

4U 1705-44 is a Galactic LMXB that exhibits X-ray bursts, and hence contains a neutron star (Langmeier et al. 1987). It exhibited a Type A PDS in the faintest of four EXOSAT observations, and a Type B PDS in the next faintest; the respective 1–11 keV fluxes were \( 1.3 \times 10^{-9} \) and \( 1.8 \times 10^{-9} \, \text{erg cm}^{-2} \text{ s}^{-1} \) (Langmeier et al. 1987 1989). Hence, an accurate distance would yield a tight constraint on \( l_c \). The distance to 4U 1705–44 has been estimated using X-ray bursts as standard candles (see Kuulkers et al. 2003); Christian & Swank (1997) find a distance of 11 kpc from Einstein data, while Cornelisse et al. (2002) obtain a distance of 8.6 kpc using data from BeppoSAX. If we assume that the distance lies between these two values, \( l_c = 0.10 \pm 0.04 \).

2.3. \( l_c \) for Galactic black hole LMXBs

Of the eighteen confirmed Galactic black hole X-ray binaries, three are high mass X-ray binaries (HMXBs), and are persistently bright, and the rest are transient LMXBs. In general outbursts last several months and the X-ray luminosity can increase by a factor of \( 10^7 \); outbursts are on average separated by years of quiescence (Chen et al. 1997; in’t Zand et al. 2004). The outbursts are hysteretic in that the transition from the low/hard state to the high/soft state during the rise of the outburst occurs at a higher luminosity than transition from the high/soft state to the low/hard state in decay (e.g. Miyamoto et al. 1995; Maccarone & Coppi 2003); hence, estimates of \( l_c \) in
black hole LMXBs were restricted to those that have been observed during the rise of the outburst, and the subset of outbursts where the transition from low/hard state to high/soft state were not made (see Brocksopp et al. 2004, for a review). Unfortunately, most X-ray observations of Galactic black hole LMXBs during outburst have been during the decay phase, and are hence unsuitable for this work.

Two of the confirmed black hole LMXBs, GX 339−4 (Zdziarski et al. 2004) and XTE J1550−564 (Rodriguez et al. 2003), have recently been caught during the rise of the outburst by the RXTE-ASM, allowing monitoring of the entire outburst by the main instruments of RXTE. Both systems exhibited spectral transitions at bolometric luminosities of ~20%. This corresponds to 0.3–10 keV luminosities of ~10% Eddington, i.e. $l_c$ ~0.1 in the 0.3–10 keV band.

Additionally, nine black hole LMXBs have exhibited outbursts where they remained in the low/hard state (for a review see Brocksopp et al. 2004). Five of these have published Type A PDS, distances and mass estimates; we have used published results to obtain the corresponding minimum value of $l_c$ in the 0.3–10 keV band. The results are presented in Fig. 3; the mass of GRS 1737−31 is given by the mass range of known black holes. We note that Esin et al. (1998) find that the SED of GRO 0422+32 at the peak of its outburst is consistent with an accretion rate of ~ $m_{\text{crit}}$; this is interesting because we find the 0.3–10 keV luminosity of GRO 0422+32 at that time to be 10±3% Eddington. Results from the other four black hole LMXBs are also consistent with $l_c$ ~0.1 in the 0.3–10 keV band, although they cannot constrain $l_c$.

2.4. Our empirical value of $l_c$

The data are all consistent with $l_c$ in the 0.3–10 keV band of 0.10±0.04 for neutron star and black hole LMXBs. Using $L_{\text{NS}} = 3.1 M_\odot$, we classify LMXBs that exhibit Type A PDS at >$4\times10^{37}$ erg s$^{-1}$ as likely black hole LMXBs. So far we have identified 11 candidates in the core of M31, including RX J0042.3+115 (Barnard et al. 2003) and CXOM31 J004303.2+411528 (Barnard, Kolb, & Osborne 2004).

3. The bright X-ray population of the central region of M31

The 0.3–10 keV luminosities and PDS of the 63 brightest X-ray sources were analysed for each of the four XMM-Newton observations of the central region of M31. The sample is likely to be dominated by X-ray binaries, and since the field of view is dominated by the bulge, they are most likely to be LMXBs. The sample was selected on the criterion that the average 0.3–10 keV EPIC-pn intensity was greater than 0.02 count s$^{-1}$ in at least one of the four observations. For each observation of each source, a PDS was made of the combined EPIC, background subtracted, 0.3–10 keV lightcurve; each PDS was averaged over many intervals, which were divided into 128 bins of 5.2 s duration. Luminosities were obtained from the unabsorbed 0.3–10 keV flux given by best fit models to EPIC-pn spectra.

From this sample, all known foreground objects and background AGN identified by Kong et al. (2002) were filtered out, as were sources that were resolved into multiple X-ray sources by Chandra. Furthermore, any observation where the classification of the PDS was ambiguous was also rejected.

Of the 167 observations accepted, 76 exhibited Type A PDS and 91 exhibited Type B or flat PDS. X-ray sources that exhibited Type A PDS at luminosities higher than $4\times10^{37}$ erg s$^{-1}$ in at least...
Fig. 4.— Cumulative distribution function (CDF) vs. 0.3–10 keV luminosity for Type A PDS for the brightest X-ray sources in four XMM-Newton observations of the central region of M31. The bottom panel shows the CDF for the whole sample, and the sample is divided into black hole candidates and non-black hole candidate in the middle and top panels respectively. The vertical lines represent the range of $L_c$ for a 1.4 $M_\odot$ neutron star primary, derived from Fig. 2. The most striking feature is the break in the CDF of the total sample, at $\sim 1.5-3 \times 10^{37}$ erg s$^{-1}$; this is what we expected for an X-ray population dominated by LMXBs with a 1.4 $M_\odot$ neutron star primary, if $L_c$ has a constant value of $\sim 0.1$. One of the four observations were classed as black hole candidates, the rest were classed as non-black holes. We note that these black hole identifications are neither certain nor complete.

Cumulative distribution function (CDFs) of Type A PDS vs luminosity are presented in Fig. 4; the CDFs for the non-black hole population is shown in the top panel; the black hole CDF is shown in the middle panel, and the CDF for the whole sample is shown in the bottom panel. The vertical lines indicate the range of transition $L_c$ obtained from the globular clusters in Fig. 1, assuming 1.4 $M_\odot$ neutron star primaries. We see that the total CDF exhibits a natural break in the region $1.5-3 \times 10^{37}$ erg s$^{-1}$, which is exactly what we would expect for a population dominated by 1.4 $M_\odot$ neutron star LMXBs, if $L_c \sim 0.1$. Also, 50% of the 51 non-BH Type A PDS were seen at luminosities below $1.0 \times 10^{37}$ erg s$^{-1}$, despite the expected strong bias towards high luminosity Type A PDS due to improved statistics.

The CDFs vs. luminosity of Type B and flat PDS for the non-black hole, black hole and total populations of X-ray sources are presented in the top, middle and bottom panels of Fig. 5 respectively. As expected, the majority ($\sim 80\%$) are observed at luminosities higher than the range of $L_c$ values obtained from the globular cluster population, and none are observed below the lower limit to $L_c$; a break occurs around $10^{38}$ erg s$^{-1}$, further indicating that most of the sources contain neutron stars. We do not expect all Type B PDS to occur above $L_c$, since some of the sources are likely to be transient, and hence hysteretic. Most significantly, none of our black hole candidates exhibit Type B PDS below $4 \times 10^{37}$ erg s$^{-1}$; this provides strong support for the idea that Type B PDS are observed at higher luminosities than Type A PDS, and suggests that these black hole candidates are not hysteretic, and are perhaps persistently bright. We note however, that only 10 Type B PDS were observed in our black hole candidates.

4. RX J0042.4+4112

RX J0042.4+4112 is located at 00h42m28.268 +41°12′22″76 (Kong et al. 2002) and is particularly exciting because, it is a black hole LMXB.
Fig. 6.— Power density spectra of 0.3–10 keV lightcurves from XMM-Newton observations of RX J0042.4+4112. The PDS are Leahy normalized, so that the Poisson noise has a power of 2.

Table 1: The spectral index (\(\alpha\)), luminosity and PDS Type of RX J0042.4+4112 in each of three XMM-Newton observations.

<table>
<thead>
<tr>
<th>Observation</th>
<th>(\Gamma)</th>
<th>(L^a)</th>
<th>PDS Type</th>
</tr>
</thead>
<tbody>
<tr>
<td>2000 Jun 25</td>
<td>1.71±0.07</td>
<td>9.1±0.6</td>
<td>FLAT (B?)</td>
</tr>
<tr>
<td>2001 Jun 9</td>
<td>1.74±0.06</td>
<td>8.0±0.5</td>
<td>A</td>
</tr>
<tr>
<td>2002 Jan 6</td>
<td>1.79±0.05</td>
<td>8.4±0.5</td>
<td>A</td>
</tr>
</tbody>
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\(a\)0.3–10 keV luminosity / 10^{37} \text{ erg s}^{-1}

by our classification that also exhibited a flat (i.e., not Type A) PDS, allowing us to estimate a primary mass. The PDS of combined EPIC (MOS1 + MOS2 + PN) 0.3–10 keV lightcurves from three XMM-Newton observations of RX J0042.4+4112 are presented in Fig. 6. The 0.3–10 keV SEDs were well described by an absorbed power law with spectral index \(\Gamma \sim 1.7–1.8\). Table 1 gives \(\Gamma\), the luminosity and PDS type for each observation.

RX J0042.4+4112 clearly exhibits Type A PDS in the 2001 and 2002 observations. However, no variability is detected in the 2000 observation, despite the higher luminosity. Since we see Type A PDS in the 2001 and 2002 observations, the lack of variability and higher luminosity in the 2000 observation is consistent with a Type B PDS. This would imply that the \(l > l_c\) in the 2000 observation. If the timing states of RX J0042.4+4112 indeed resemble those of Galactic LMXBs, then we can estimate the mass of the primary.

We find that \(L_c = 8.8 \pm 0.9 \times 10^{37} \text{ erg s}^{-1}\); assuming that \(l_c = 0.10 \pm 0.04\), the primary mass is \(7 \pm 2 \text{ M}_\odot\), which is above the maximum mass of a neutron star.

RX J0042.4+4112 is consistently bright in the 4 XMM-Newton observations, as well as in the 1991 and 1992 ROSAT surveys of M31 (Supper et al. 1997, 2001), the 1979 Einstein observation (Trinchieri & Fabbiano 1991), and in all 30 Chandra observations of the source made between October 1999 and September 2002. Hence RX J0042.4+4112 is likely to have been persistently bright for over 20 years.

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REFERENCES


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